

## OBSERVATIONAL ASTRONOMY

Astronomy is fundamentally an observational science, and as such it is important for astronomers and astrophysicists to understand how their data are collected and analyzed. This book is a comprehensive review of current observational techniques and instruments.

Featuring instruments such as Spitzer, Herschel, Fermi, ALMA, Super-Kamiokande, SNO, IceCube, the Auger Observatory, LIGO, and LISA, the book discusses the capabilities and limitations of different types of instruments. It explores the sources and types of noise and provides statistical tools necessary for interpreting observational data. Due to the increasingly important role of statistical analysis, the techniques of Bayesian analysis are discussed, along with sampling techniques and model comparison.

With topics ranging from fundamental subjects such as optics, photometry, and spectroscopy, to neutrinos, cosmic rays, and gravitational waves, this book is essential for graduate students in astronomy and astrophysics.

EDMUND C. SUTTON is Associate Professor in the Astronomy Department at the University of Illinois. His research has been primarily in infrared and submillimeter astronomy with an emphasis on instrumentation.

# OBSERVATIONAL ASTRONOMY

## Techniques and Instrumentation

EDMUND C. SUTTON

*University of Illinois*



Cambridge University Press  
978-1-107-01046-8 — Observational Astronomy  
Edmund C. Sutton  
Frontmatter  
[More Information](#)

CAMBRIDGE UNIVERSITY PRESS  
Cambridge, New York, Melbourne, Madrid, Cape Town,  
Singapore, São Paulo, Delhi, Tokyo, Mexico City

Cambridge University Press  
The Edinburgh Building, Cambridge CB2 8RU, UK

Published in the United States of America by Cambridge University Press, New York

[www.cambridge.org](http://www.cambridge.org)  
Information on this title: [www.cambridge.org/9781107010468](http://www.cambridge.org/9781107010468)

© E. C. Sutton 2012

This publication is in copyright. Subject to statutory exception  
and to the provisions of relevant collective licensing agreements,  
no reproduction of any part may take place without the written  
permission of Cambridge University Press.

First published 2012

Printed in the United Kingdom at the University Press, Cambridge

*A catalog record for this publication is available from the British Library*

*Library of Congress Cataloging in Publication data*  
Sutton, Edmund Charles.

Observational astronomy : techniques and instrumentation / Edmund C. Sutton.  
p. cm.

ISBN 978-1-107-01046-8 (hardback)

1. Astronomy – Textbooks. 2. Astronomy – Observations – Textbooks. I. Title.  
QB43.3.S88 2011  
520–dc23  
2011030682

ISBN 978-1-107-01046-8 Hardback

Additional resources for this publication at [www.cambridge.org/9781107010468](http://www.cambridge.org/9781107010468).

Cambridge University Press has no responsibility for the persistence or  
accuracy of URLs for external or third-party internet websites referred to  
in this publication, and does not guarantee that any content on such  
websites is, or will remain, accurate or appropriate.

## Contents

<i>List of illustrations</i>	<i>page</i> xiv
<i>List of tables</i>	xx
<i>Preface</i>	xxi
<i>Acknowledgements</i>	xxiii
<b>1 Astrophysical information</b>	<b>1</b>
1.1 Electromagnetic radiation	1
1.2 Other carriers of information	2
1.3 Intervening regions	3
1.3.1 Intergalactic/interstellar medium	3
1.3.2 Interplanetary medium	5
1.3.3 Earth's atmosphere	5
Exercises	12
<b>2 Photometry</b>	<b>14</b>
2.1 Specific intensity (brightness)	14
2.2 Étendue	14
2.3 Moments of the specific intensity	16
2.4 Energy density	17
2.5 Flux from a surface of uniform brightness	17
2.6 Blackbody radiation	18
2.7 Atmospheric extinction (calibration)	20
2.8 Absolute calibration	22
2.9 Photometric magnitudes	23
Exercises	25
<b>3 Positional astronomy</b>	<b>28</b>
3.1 Fundamental reference system	28
3.2 Time systems	28
3.2.1 Atomic time	28

	3.2.2	Astronomical time scales	30
	3.2.3	Sidereal time	31
	3.2.4	Solar time	31
3.3		Spherical astronomy	32
	3.3.1	Spherical coordinates (in general)	32
	3.3.2	Latitude and longitude	33
	3.3.3	Equatorial coordinates	33
	3.3.4	Horizon coordinate system (alt/az)	34
	3.3.5	Conversion formulae (alt/az $\leftrightarrow$ ha/dec)	35
	3.3.6	Ecliptic coordinates	35
	3.3.7	Galactic coordinates	36
	3.3.8	Spherical trigonometry	37
	3.3.9	Rotation matrices	37
3.4		Epoch	39
3.5		Changes in equatorial coordinates	40
	3.5.1	Proper motion	40
	3.5.2	Precession	41
	3.5.3	Nutation	41
	3.5.4	Parallax	42
	3.5.5	Aberration of starlight	42
	3.5.6	Reduction of celestial coordinates (overview)	43
	3.5.7	Gravitational deflection of light	44
	3.5.8	Refraction	45
	3.5.9	Parallactic angle	46
3.6		Astrometry	46
	3.6.1	Historical techniques	46
	3.6.2	Hipparcos	48
		Exercises	49
<b>4</b>		<b>Fourier transforms</b>	<b>52</b>
	4.1	Fourier series	52
	4.2	Fourier integrals	54
	4.2.1	Relationship to the Dirac delta (impulse) function	55
	4.2.2	Parseval's theorem (Rayleigh's theorem)	56
	4.2.3	Properties of Fourier transforms	57
	4.2.4	Convolution	58
	4.2.5	Autocorrelation (Wiener–Khinchin theorem)	58
	4.2.6	Common functions and Fourier transform pairs	59
	4.2.7	Aliasing and Shannon's sampling theorem	60

*Contents*

vii

	4.3 Higher-dimensional Fourier transforms	62
	4.3.1 Hankel (Fourier–Bessel) transforms	62
	Exercises	63
<b>5</b>	<b>Detection systems</b>	<b>66</b>
	5.1 Interaction of radiation and matter	66
	5.2 Photoelectric effect	66
	5.3 Compton scattering	67
	5.4 Pair production	69
	5.5 Electromagnetic wave interactions	70
	5.6 Optical and ultraviolet detectors	70
	5.6.1 Photomultipliers	70
	5.6.2 Other electron multiplication devices	72
	5.6.3 Solid state detectors	74
	5.7 Infrared astronomy	78
	5.7.1 Infrared photoconductors	80
	5.7.2 NICMOS	82
	5.7.3 Bolometers	83
	5.7.4 Spitzer	84
	5.7.5 Herschel	85
	5.7.6 WFIRST	85
	Exercises	86
<b>6</b>	<b>Orthodox statistics</b>	<b>87</b>
	6.1 Probability distributions	87
	6.1.1 Binomial distribution	88
	6.1.2 Poisson distribution	89
	6.1.3 Gaussian (normal) distribution	91
	6.2 Moments of a probability distribution	92
	6.3 Characteristic (moment-generating) function	92
	6.4 Central limit theorem	94
	6.5 Experimental data	95
	6.6 Chi-squared ( $\chi^2$ ) distribution	97
	6.7 Student's t-distribution	99
	6.8 Robust estimation	100
	6.9 Propagation of errors	102
	Exercises	104
<b>7</b>	<b>Stochastic processes and noise</b>	<b>105</b>
	7.1 Stochastic process	105
	7.1.1 Stationary process	106
	7.2 Spectral density of a Poisson random process	106
	7.3 Spectral density of a Gaussian random process	108

viii	<i>Contents</i>	
7.4	The transformation $y = x^2$	109
7.5	Filtering	110
7.5.1	Low pass filtering	110
7.6	Estimation in the presence of Gaussian noise	112
7.7	Photon noise	112
7.8	Thermal noise	113
	Exercises	115
<b>8</b>	<b>Optics</b>	<b>117</b>
8.1	Geometrical optics	117
8.1.1	Paraxial optics (a first order theory)	119
8.1.2	Seidel aberrations (a third order theory)	124
8.1.3	Higher order terms	129
8.1.4	Telescope design	130
8.1.5	Other aspects of telescope design	132
8.1.6	Gravitational lensing	133
8.2	Dispersion	134
8.2.1	Origin of the refractive index	134
8.2.2	Fresnel coefficients	135
8.3	Physical optics	138
8.3.1	Vector and scalar diffraction	139
8.3.2	Kirchhoff diffraction theory	139
8.3.3	Fresnel and Fraunhofer approximations	143
8.3.4	Diffraction with aberrations	149
8.4	Imaging	150
8.5	Addendum	153
	Exercises	155
<b>9</b>	<b>Interference</b>	<b>158</b>
9.1	Mutual coherence function and complex degree of coherence	158
9.2	Quasi-monochromatic radiation	158
9.3	Young's two-slit experiment	160
9.4	Michelson interferometer	162
9.5	Michelson stellar interferometer	163
9.6	Van Cittert–Zernike theorem	164
9.7	Étendue of coherence	166
9.7.1	One approach	166
9.7.2	An alternate approach	167
9.8	Aperture synthesis	168
9.8.1	Arrays of antennas	168
9.9	Caveat	170

*Contents*

ix

9.10	Fourth order coherence	170
9.10.1	Intensity interferometry	171
	Exercises	172
<b>10</b>	<b>Spectroscopy</b>	<b>173</b>
10.1	Multiple beam interference	173
10.1.1	Airy function	174
10.1.2	Anti-reflection coating	175
10.1.3	Enhanced reflection coating	176
10.1.4	Interference filters	177
10.2	Fabry–Perot interferometer (etalon)	178
10.3	Fourier transform spectrometer	180
10.4	Prism spectrograph	181
10.4.1	Prism applications	182
10.5	Diffraction gratings	183
10.5.1	Grating properties	184
10.5.2	Grating profiles	185
10.5.3	Czerny–Turner spectrograph	185
10.5.4	Echelle spectrograph	186
10.5.5	Grism spectroscopy	187
10.5.6	Fiber optic spectroscopy	187
	Exercises	188
<b>11</b>	<b>Ultraviolet, x-ray, and gamma ray astronomy</b>	<b>190</b>
11.1	Telescopes and imaging	190
11.1.1	X-ray telescopes	190
11.1.2	Collimators	192
11.1.3	Tracking designs	192
11.1.4	Coded apertures	192
11.2	Detectors	194
11.2.1	Proportional counters	194
11.2.2	Solid state detectors	195
11.2.3	Scintillators	195
11.2.4	Spark chambers	196
11.3	Recent missions	196
11.3.1	ROSAT	196
11.3.2	Compton Gamma Ray Observatory	197
11.3.3	Extreme Ultraviolet Explorer	199
11.3.4	ASCA	200
11.3.5	Rossi X-ray Timing Explorer	200
11.3.6	BeppoSAX	201
11.3.7	FUSE	202



11.3.8	Chandra	202
11.3.9	XMM-Newton	203
11.3.10	INTEGRAL	205
11.3.11	GALEX	206
11.3.12	Swift	206
11.3.13	Fermi gamma ray space telescope	207
11.4	Possible future missions	207
11.4.1	IXO	207
11.4.2	MAXIM or BHI	207
	Exercises	208
<b>12</b>	<b>Radio receivers, spectrometers, and interferometers</b>	<b>209</b>
12.1	Astrophysical radio sources	209
12.2	Fundamentals of radio receivers	209
12.2.1	Linear systems	210
12.2.2	Quantum noise limit	211
12.2.3	Components in series	211
12.2.4	Low noise GaAs FET amplifiers	212
12.2.5	Radio frequency mixers	214
12.2.6	Detectors and the radiometer equation	217
12.3	Precision radiometry of the CMB	218
12.3.1	COBE	218
12.3.2	WMAP	218
12.3.3	Planck	220
12.3.4	Atacama Cosmology Telescope	220
12.4	Radio spectrometers	221
12.4.1	Autocorrelation spectrometers	221
12.4.2	Filter banks	222
12.4.3	Acousto-optical spectrometers	222
12.5	Radio antennas	223
12.5.1	Antenna patterns	223
12.5.2	Antenna temperature	225
12.5.3	Special antenna designs	226
12.6	Radio interferometry	226
12.6.1	Basic two-element interferometer	226
12.6.2	Interferometer arrays	228
12.6.3	Correlators	229
12.6.4	Fourier inversion	230
<b>13</b>	<b>Modern statistical methods</b>	<b>233</b>
13.1	Bayes' theorem	233
13.2	Maximum likelihood	235

<i>Contents</i>		xi
13.3	So what is Bayesian inference?	236
13.3.1	Example 1	236
13.3.2	Example 2	237
13.4	Maximum entropy	239
13.5	Uninformative priors	241
13.5.1	Location priors	241
13.5.2	Scale priors	241
13.5.3	Positive, additive distributions	241
13.6	Inverse problems	242
13.7	Sampling the posterior	243
13.7.1	Rejection sampling	244
13.7.2	Metropolis–Hastings algorithm	245
13.7.3	Gibbs sampling	246
13.7.4	Mixing behavior	247
13.8	Model comparison	247
13.9	Malmquist (truncation) bias	250
13.10	Censoring	251
13.11	Confidence limits	253
	Exercises	255
<b>14</b>	<b>Neutrino detectors</b>	<b>257</b>
14.1	Neutrinos	257
14.2	Solar neutrino production	258
14.3	Supernova production	260
14.4	Atmospheric neutrinos	260
14.5	Neutrino oscillations	262
14.5.1	Vacuum oscillations	262
14.5.2	Matter oscillations	264
14.5.3	Conclusions	264
14.6	Radiochemical (transmutational) detectors	265
14.6.1	Chlorine	265
14.6.2	Gallium	267
14.6.3	Other targets	268
14.7	Čerenkov detectors	269
14.7.1	Kamiokande and Super-Kamiokande	271
14.7.2	Sudbury Neutrino Observatory	276
14.7.3	IceCube	279
14.8	Scintillation detectors: Borexino	282
14.9	Cosmological implications	283
14.10	Background of supernova neutrinos	283
	Exercises	284

<b>15</b>	<b>Cosmic ray detectors</b>	<b>285</b>
15.1	Properties of cosmic rays	285
15.2	Intervening regions	287
15.2.1	Magnetic fields	287
15.2.2	Spallation reactions	289
15.2.3	Interstellar ionization losses	290
15.2.4	Bremsstrahlung	290
15.2.5	Synchrotron losses	292
15.2.6	Inverse Compton losses	293
15.2.7	Pair production	294
15.2.8	GZK effect	294
15.2.9	Decays	295
15.2.10	Atmospheric interactions	295
15.3	Detectors	295
15.3.1	Ionization detectors	296
15.3.2	Bremsstrahlung	297
15.3.3	Čerenkov radiation	300
15.3.4	Transition radiation	300
15.4	Balloon-borne and spacecraft missions	301
15.4.1	1990s and early 2000s	301
15.4.2	TRACER and CREAM	301
15.4.3	PAMELA	303
15.4.4	Alpha Magnetic Spectrometer	303
15.5	Extensive air showers	303
15.5.1	High Resolution Fly's Eye	307
15.5.2	Pierre Auger Observatory	309
15.5.3	Telescope Array (TA) project	310
15.5.4	Atmospheric Čerenkov Telescope Array	311
15.5.5	JEM-EUSO	311
15.6	Particle acceleration	312
	Exercises	313
<b>16</b>	<b>Gravitational waves</b>	<b>314</b>
16.1	Characteristics of gravitational radiation	314
16.2	Sources of gravitational waves	316
16.3	Ground-based interferometric detectors	320
16.3.1	Fabry–Perot	323
16.3.2	Recycling interferometers	325
16.3.3	Lasers	325
16.3.4	Seismic noise	326

<i>Contents</i>		xiii
16.3.5	Quantum limit, shot noise, and radiation pressure fluctuations	328
16.3.6	Thermal noise	330
16.3.7	Other factors	330
16.3.8	Performance	331
16.3.9	Squeezed states	332
16.4	Space-based interferometric detectors	333
16.5	Other systems	336
16.6	Data analysis	336
	Exercises	337
<b>17</b>	<b>Polarimetry</b>	<b>340</b>
17.1	Sources of polarized radiation	340
17.1.1	Synchrotron radiation	340
17.1.2	Zeeman effect	340
17.1.3	Thermal emission	343
17.1.4	Scattering	343
17.1.5	Primordial polarization	344
17.2	Propagation effects	345
17.3	Polarization-sensitive devices	345
17.4	Analysis of polarization states	346
17.4.1	Stokes parameters	346
17.4.2	Mueller matrices	348
17.4.3	Jones vectors and matrices	349
17.5	Polarization measurement	350
17.5.1	Analysis of weak field splittings	351
17.6	Optical polarimetry	351
17.7	Radio polarimetry and calibration	352
	Exercises	353
	<i>Appendix A</i> Physical constants and units	355
	<i>Appendix B</i> Acronyms	356
	<i>Appendix C</i> Additional reading	363
	<i>References</i>	369
	<i>Index</i>	378

The color plates will be found between pages 232 and 233.

## Illustrations

1.1	Atmospheric transmission from 10 to 1000 GHz	<i>page</i> 6
1.2	Atmospheric transmission from 480 to 500 GHz	7
1.3	Scattering from bound electrons	10
1.4	Seeing and scintillation	11
2.1	Geometry defining the specific intensity	15
2.2	Conservation of étendue	15
2.3	Changes in image size and solid angle	15
2.4	Flux through an aperture	16
2.5	Relation between energy density and mean intensity	17
2.6	Integrated flux from a surface of uniform brightness	18
2.7	Opacity of Earth's atmosphere	21
2.8	Sky-dip method for calculating zenith opacity	22
2.9	Sidelobes of a radio telescope	22
2.10	Filter response of standard photometric systems	24
3.1	Cesium clock	29
3.2	Sidereal time	31
3.3	Spherical coordinates	32
3.4	Latitude and longitude on Earth	33
3.5	Fixed and rotating equatorial coordinate systems	34
3.6	Rotating and fixed meridians	34
3.7	Horizon coordinate system	35
3.8	Ecliptic coordinate system	36
3.9	Galactic coordinate system	36
3.10	Relationship between galactic and equatorial coordinates	38
3.11	Spherical trigonometry	39
3.12	Rotation matrices	39
3.13	Precession	41
3.14	Parallax $\Pi$ for a star at distance $D$	42
3.15	Aberration of starlight	43
3.16	Gravitational deflection of light at the limb of the Sun	44
3.17	Refraction in a plane-parallel atmosphere	45
3.18	Wide field astrometric errors	47

*List of illustrations*

xv

3.19	Altitude/azimuth coordinate system	49
4.1	Square wave as a finite Fourier series	54
4.2	An exponentially decaying wave and its power spectrum	56
4.3	Aliasing	60
4.4	Sampling theorem	61
4.5	Gibbs' phenomenon	64
5.1	X-ray total absorption cross section for Xe	67
5.2	Compton scattering	68
5.3	Compton cross sections	69
5.4	Pair production cross section in lead	70
5.5	Schematic view of a photomultiplier	71
5.6	A Multi-Anode Microchannel Array (MAMA)	73
5.7	Silicon with donors and acceptors	74
5.8	PN junction	75
5.9	Metal-oxide-semiconductor (MOS) device	77
5.10	CCD readouts	78
5.11	Thermal backgrounds for infrared telescopes	79
5.12	Intrinsic and extrinsic photoconductors	80
5.13	Photoconductive gain	81
5.14	Thermal circuit for a bolometer	83
6.1	Probability density and cumulative probability	88
6.2	Binomial distribution	89
6.3	Poisson distributions for $a = 1, 1.5,$ and $2$	90
6.4	Poisson statistics based on binomial distribution	90
6.5	Gaussian distribution	92
6.6	Convergence in central limit theorem	95
6.7	Rainfall statistics for Urbana, Illinois	96
6.8	Probability density function for Student's t-distribution	101
6.9	Outliers and robust estimation	101
6.10	Outliers and least-squares fitting	102
7.1	Autocorrelation of a Poisson random process	107
7.2	The transformation $y = x^2$	109
7.3	Filtering	110
7.4	Ideal low pass filter	111
7.5	Thermal noise passed by a low pass filter	114
7.6	Johnson noise	114
8.1	Law of reflection	118
8.2	Snell's law	119
8.3	Paraxial optics and refraction	120
8.4	Paraxial optics and reflection	121
8.5	Transverse and angular magnification	122
8.6	Thin and thick lenses	123
8.7	Spherical aberration for a mirror	125
8.8	Spherical aberration for a lens	126
8.9	Coma	126
8.10	Coma in the tangential and sagittal planes	127

xvi	<i>List of illustrations</i>	
8.11	Astigmatism	128
8.12	Astigmatism in the tangential and sagittal planes	128
8.13	Pincushion and barrel distortion	129
8.14	Petzval field curvature	130
8.15	Classical Cassegrain telescope	130
8.16	Schmidt telescope	132
8.17	Drude–Lorentz model of the refractive index	134
8.18	Dispersion relation	136
8.19	Fresnel coefficients	136
8.20	S-polarization	137
8.21	Application of Green’s theorem to Kirchhoff diffraction	140
8.22	Kirchhoff diffraction for an opaque screen with a hole	141
8.23	Huygens–Fresnel principle	142
8.24	Geometry of the Poisson spot	142
8.25	Intensity distribution along the axis for the Poisson spot	143
8.26	Geometry for the Fresnel approximation	144
8.27	Cylindrical coordinates for diffraction from a circular aperture	145
8.28	The Airy pattern: diffraction from a circular aperture	146
8.29	Diffraction by a straight edge	147
8.30	The Cornu spiral: the Fresnel integrals $S(u)$ and $C(u)$	148
8.31	Diffraction pattern for a point source and a straight edge	149
8.32	Point spread function	150
8.33	Modulation transfer function	151
8.34	A rectangular pixel	153
8.35	Nyquist sampling with a grid of pixels	154
8.36	Marginal and paraxial rays	155
9.1	Electric fields separated in time and space	159
9.2	Quasi-monochromatic radiation	159
9.3	Gaussian power spectral density and autocorrelation functions	159
9.4	Coherence length of quasi-monochromatic radiation	160
9.5	Young’s two-slit experiment	161
9.6	Fully modulated fringes	161
9.7	Loss of coherence	162
9.8	Michelson interferometer	162
9.9	Michelson stellar interferometer	163
9.10	Fringes from two point sources being washed out	164
9.11	Coherence of radiation	165
9.12	Radiation pattern of a radio telescope	167
9.13	Radio interferometer	169
9.14	Linear array of dipoles	169
9.15	Antenna pattern of a linear array	170
10.1	Interference between plane-parallel dielectric interfaces	174
10.2	Airy function	175
10.3	Anti-reflection coating	176
10.4	Enhanced reflectivity coating	177
10.5	Fabry–Perot interferometer	178

*List of illustrations*

xvii

10.6	Fabry–Perot interferometers in series	179
10.7	Fourier transform spectrometer	180
10.8	Angular deflection by a prism	181
10.9	Prism spectrograph	182
10.10	Objective prism	183
10.11	Principle of a grating	184
10.12	Blazed grating	185
10.13	Czerny–Turner spectrograph	186
10.14	Echelle grating	186
10.15	Grism	187
11.1	Examples of total internal reflection	191
11.2	Wolter type-I telescope with nested optics	191
11.3	Illustration of a pinhole camera and a coded aperture mask	193
11.4	Photoionization within a gas proportional counter	194
11.5	Gas proportional counter	195
11.6	COMPTEL detector on CGRO	199
11.7	EGRET detector on CGRO	200
11.8	Proportional counter unit from RXTE PCA	201
11.9	Chandra ACIS imaging and spectroscopy arrays	203
11.10	Chandra HRC	204
11.11	Coded aperture masks for SPI and IBIS	205
12.1	Thermal bremsstrahlung and synchrotron emission	210
12.2	Examples of linear systems	210
12.3	Amplifiers connected in series	212
12.4	FET amplifier	213
12.5	Mixer frequencies	215
12.6	Schottky diode	215
12.7	SIS tri-layer	216
12.8	Simple radiometer system	217
12.9	WMAP results	219
12.10	Polarization-sensitive bolometers	220
12.11	Autocorrelation spectrometer	221
12.12	Bragg reflection	222
12.13	Acousto-optical spectrometer	223
12.14	Ruze theory	224
12.15	Radio antenna pattern	224
12.16	Radio interferometer	227
12.17	Baseline variation due to Earth’s rotation	228
13.1	Probability of a neutrino detection	238
13.2	Bayesian posterior	239
13.3	Example of rejection sampling	244
13.4	Metropolis algorithm	246
13.5	Model comparison	248
13.6	Evidence comparison	249
13.7	Example of Malmquist bias	251
13.8	Example of censoring	252



14.1	SSM neutrino flux predictions	259
14.2	Atmospheric neutrinos	261
14.3	System of two coupled pendulums	263
14.4	Neutrino mixing parameters	265
14.5	Ray Davis during construction of his neutrino detector	266
14.6	DUSEL laboratory	269
14.7	Neutrino–electron elastic scattering	270
14.8	Geometry of Čerenkov radiation	270
14.9	Inner portion of Super-K detector	272
14.10	Outer portion of Super-K detector	272
14.11	Super-K event topology	273
14.12	Electron-like and muon-like Super-K events	274
14.13	Neutrino cross sections for light and heavy water	275
14.14	Creighton mine	276
14.15	Artist’s conception of SNO	277
14.16	SNO detector during construction	277
14.17	Solar neutrino fluxes	278
14.18	Photomultiplier deployed into antarctic ice	279
14.19	IceCube photomultiplier strings	280
14.20	Simulated muon event in IceCube	281
15.1	Cosmic ray spectrum	286
15.2	Cosmic ray deflection	288
15.3	Cosmic ray components	291
15.4	Electron impact parameter	292
15.5	Ionization energy losses	297
15.6	Energy loss for light nuclei	298
15.7	Isotope mass resolution	299
15.8	Ionization tracks in plastic	299
15.9	TRACER	302
15.10	$X_{\max}$	305
15.11	Shower particles arrive at ground array stations	305
15.12	HiRes air shower event	308
15.13	GZK cutoff	308
15.14	Pierre Auger Observatory	309
16.1	Linear polarization states of gravitational waves	316
16.2	Possible sources of gravitational radiation	318
16.3	Chirp signal	319
16.4	Images of interferometers	322
16.5	Locations and orientations of interferometers	323
16.6	LIGO optical layout	324
16.7	Mode cleaner	326
16.8	Quadruple pendulum prototype	327
16.9	LIGO sensitivity	331
16.10	Sensitivity goals for Einstein Telescope	332
16.11	Configuration of LISA satellites	333
16.12	Estimated LISA sensitivity	334

<i>List of illustrations</i>		xix
16.13	Possible configuration of LISA spacecraft	335
17.1	Polarization of synchrotron radiation	341
17.2	Longitudinal and transverse Zeeman effect	341
17.3	Hyperfine splitting of OH	342
17.4	Zeeman doublet and triplet showing $\pi$ and $\sigma^{\pm}$ components	343
17.5	Scattering of unpolarized light	344
17.6	Wollaston prism	346
17.7	Right circular polarization	348

## Tables

1.1	Characteristic photon energies and temperatures	<i>page</i> 2
2.1	Johnson–Cousins–Glass photometric system	25
2.2	SDSS bands	25
3.1	Hipparcos precision	48
4.1	Symmetry properties of Fourier transform pairs	55
4.2	Properties of Fourier transforms	57
4.3	Symbols for common functions	59
4.4	Fourier transform pairs	59
6.1	Reduced chi-squared	98
6.2	Chi-squared	99
6.3	Student’s t-distribution	100
13.1	One-dimensional confidence limits	253
13.2	Two-dimensional confidence limits	254
15.1	Characteristic magnetic fields and gyroradii	288
15.2	Radioactive isotopes replenished from cosmic rays	290
16.1	Gravitational wave interferometers	320
A.1	Physical constants	355
A.2	Other units	355

## Preface

This book is based on a required course for graduate students in Astronomy which I taught for a number of years at the University of Illinois. The premise of the course is that both theoretical astronomers and observers should have a basic understanding of the techniques of observational astronomy. The emphasis is on the underlying physics of the methods of detection and analytical tools (statistical and otherwise) that astronomers find useful. The great variety of current instruments and the rapid introduction of new instruments preclude an in-depth treatment of the peculiarities and idiosyncrasies of many instruments. But every instrument has its own idiosyncrasies and its own ways of corrupting the data and deceiving the observer. The topics in this book, I believe, cover the minimum which is required of anyone attempting to understand or interpret observational astronomy data.

Throughout the book equations are given in mks (SI) units so that it is easy to relate the discussion to practical quantities such as volts and watts. This is true even in the chapter on gravitational waves, a subject for which many texts and references use geometrized units ( $c = 1$ ,  $G = 1$ ). I prefer to keep  $c$  and  $G$  around rather than having to figure out where to put them when I need to calculate power. I also like being able to check equations using dimensional analysis. In the text other units are freely worked in. Among astronomers, the gauss remains firmly fixed as the unit of magnetic flux density. And astronomers frequently use other cgs units. For example, cross sections are always in  $\text{cm}^2$ . And of course there is a plethora of astronomical units such as pc, AU(!), and  $M_{\odot}$ . An appendix is provided with physical constants in both mks and cgs units and with a list of other units used and their equivalents in mks and cgs units.

The reader will note that the chapters on neutrinos, cosmic rays, and gravitational waves are of a different nature than other parts of the book. These fields are sufficiently specialized that it is difficult to separate purely observational issues from the underlying science. Therefore, in these chapters I freely go back and forth between design and scientific goals.

In addition to the color plates, there are color versions of a large number of other figures. The complete set of color figures may be accessed and/or downloaded through this book's website: [www.cambridge.org/9781107010468](http://www.cambridge.org/9781107010468).

I am well aware of other topics that I could have included in this book. In particular, I regret not being able to include a thorough discussion of adaptive optics and not covering topics in astroparticle physics.

The outlook for possible future instruments has changed markedly since much of this text was written, largely due to budgetary constraints. A funding increment for DUSEL (Chapter 14) by the National Science Foundation was recently rejected by the US National Science Board. The fate of DUSEL currently rests with its remaining US sponsor, the Department of Energy. WFIRST (Chapter 5) remains a high priority project for NASA. If ESA assigns a similarly high priority to its Euclid mission, a merger of these projects is likely to be considered. The US commitments to IXO (Chapter 11) and LISA (Chapter 16) are very much in doubt. These international collaborations are expected to continue, but reduced financial support could lead to delays and reductions in scope. In any event, these instrument concepts are the current state of the art. Astronomers constantly need to readjust their plans in light of financial realities. If better ways can be found to pursue some of these scientific objectives, now is certainly the time for them.

## Acknowledgements

I appreciate the willingness of my colleagues Brian Fields, Athol Kemball, Ben Wandelt, and Dick Crutcher to review limited sections of this text. Their comments have been very useful. Any remaining mistakes are, of course, solely my responsibility. I encourage anyone who discovers errors of any sort to communicate them to me at [ecsutton@illinois.edu](mailto:ecsutton@illinois.edu). The publisher and I will work together to maintain an online list of any errata.

The graduate students to whom I have had the pleasure to teach this material over the years are Scott Bain, Ian Barton, Yohann Beda, Jana Bilikova, Mark Butala, Karen Camarda, Christine Cecala, Nachiketa Chakraborty, Ray Chen, Rosie Chen, Yun Chen, Hsin-Fang Chiang, Samuel Crawford, Conley Ditsworth, Joshua Dolence, Bryan Dunne, Rich Frazin, Khurram Gillani, Daniel Goscha, Philip Grathoff, Michelle Griffin, Xiaoyue Guan, Troy Hacker, Thomasanna Hail, Nicholas Hakobian, Hassan Halataei, Brett Hayes, Nathan Hearn, Nicholas Indriolo, Rishi Khatri, Soyoung Kim, Robert Klinger, Scott Kruger, Hsin-Lun Kuo, Woojin Kwon, Shih-Ping Lai, I-Jen Lee, PoKin Leung, Amy Lien, Wen-Ching Lin, Jiayi Liu, Sheng-Yuan Liu, Justin Lowry, Zarija Lukic, Britt Lundgren, Patrick Lynch, Modhurita Mitra, Rosa Murphy, Erik Nelson, Christopher Neyman, Chenping Ni, Lisa Norton, Brian O'Neill, Kuo-Chuan Pan, Vasiliki Pavlidou, Tijana Prodanovic, Ramprasad Rao, David Rebolledo Lara, Ashley Ross, Jonathan Seale, Jerry Shaw, Hotaka Shiokawa, Jeeseon Song, Thomas Spinka, Ranjani Srinivasan, Ian Stephens, Shweta Sundararajan, Konstatinos Tassis, Daniel Thayer, Glenn Thurman, Toshiya Ueta, Scott Walker, Li-Bang Wang, Shiya Wang, Yiran Wang, Rui Xue, Amit Yadav, Chao-Chin Yang, Hsiang-Yi Yang, Jeong Yim, Alfredo Zenteno, and Jie Zou. They have done much to determine the direction my course has taken and have been invaluable at finding mistakes and inconsistencies in the notes. I am indebted to them.

I appreciate the help of the production staff at Cambridge University Press, especially Claire Poole, Vince Higgs, and Abigail Jones. Ms. Sehar Tahir was helpful with TeX support. Margaret Patterson was an excellent copy editor.

I am indebted to the developers of Inkscape, which made the production of the figures relatively painless. A portion of the author's proceeds from this publication has been donated in advance to the Software Freedom Conservancy to help defray further development costs for the Inkscape Project.

Every effort has been made to acknowledge and obtain permission for all figures used in this work. Figure 2.10 is included by kind permission of R. A. Jansen. Tables 4.1, 4.2, and 4.3 are reprinted with permission from Bracewell, *The Fourier transform and its applications*, 3rd edn. ©2000 The McGrawHill Companies, Inc. Figure 5.6 is reprinted with permission from Timothy (1983) *PASP*, **95**, 573, ©1983 University of Chicago Press. The image shown in Figure 10.10 was kindly obtained by D. Ketelsen, who has granted permission for its use here. Figures 11.6, 11.7, and 11.8 reproduced by permission of the AAS. The quotation from Bahcall & Ostriker, *Unsolved problems in astrophysics* ©1997 is used by permission of Princeton University Press. Figure 14.1 reproduced by permission of the AAS. Figures 14.4, 14.12, and 14.17 are reprinted with permission from Aharmim *et al.* (2005) *Phys. Rev. C*, **72**, 055502, Ashie *et al.* (2005) *Phys. Rev. D*, **71**, 112005, and Abbasi *et al.* (2008) *Phys. Rev. Lett.*, **100**, 101101, ©2005 and 2008 by the American Physical Society. Figures 14.9 and 14.10 reprinted with permission, ©Kamioka Observatory, ICRR, University of Tokyo. Figures 14.14, 14.15, and 14.6 are courtesy of SNO. Figures 14.18, 14.19, and 14.20 are based upon work supported by the National Science Foundation under Grant Nos. OPP-9980474 (AMANDA) and OPP-0236449 (IceCube), University of Wisconsin-Madison. Figure 15.1 reprinted with permission of Annual Reviews from Beatty & Westerhoff (2009) *ARNPS*, **59**, 319; permission conveyed through Copyright Clearance Center, Inc. Figure 15.3 reprinted from Gaisser & Stanev (2008) *Phys. Lett. B*, **667**, 254 with permission from Elsevier and D. Muller. Figure 15.7a reprinted from Stone *et al.* (1998) *Sp. Sci. Rev.*, **86**, 357 with permission from Elsevier. Figure 15.7b reprinted by permission of R. Ogliore. Figure 15.9 reproduced by permission of the AAS. Figure 15.12 reprinted from Boyer *et al.* (2002) *NIMPR A*, **482**, 457 with permission from Elsevier. Figure 15.13 reprinted by permission from Abbasi *et al.* (2008) *Phys. Rev. Lett.*, **100**, 101101 ©2008 by the American Physical Society. Figure 15.14 reprinted by permission of P. Mantsch on behalf of the Auger Project. Figure 16.6 reprinted from Abadie *et al.* (2010c) *NIMPR A*, **624**, 223 with permission from Elsevier. Figure 16.8 reprinted by permission, ©Science and Technology Facilities Council and Brett Shapiro/LIGO Laboratory. Figure 16.10 reprinted from Hild *et al.* (2010) *Class. Quantum Grav.*, **27**, 15003 with permission from IoP Publishing.

And finally I thank Jean for all of her support during the writing and production of this book.